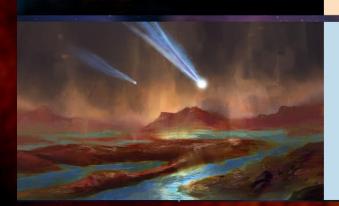
Tracing Planet Formation in the Hidden Far Infrared

Fingerprinting Planetary Reservoirs: Determine how planets form in disks around young stars, and explain the observed diversity of planets.

Are we alone?



Tracing Water to Rocky Planets: Determine the source of water in planetforming disks, and explain how water accumulates into oceans.

How did we get here?



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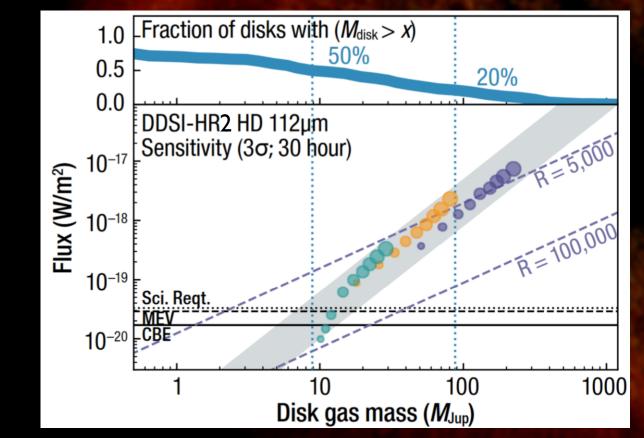


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HD will establish accurate disk masses

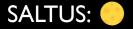


Disk to stellar mass ratios of 1% (purple), 0.4% (orange), and 0.1% (green) for 0.3 – 2.0 solar masses 56μm and 112μm lines
50% of disks can be detected at sensitivities of <10⁻¹⁹ at R=10⁵

 At lower sensitivities and R, only detect HD around most massive disks

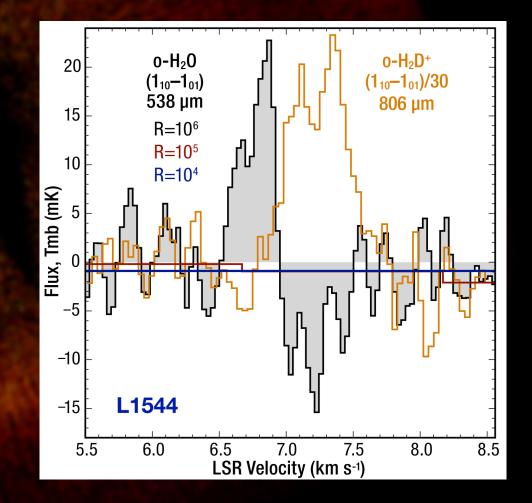






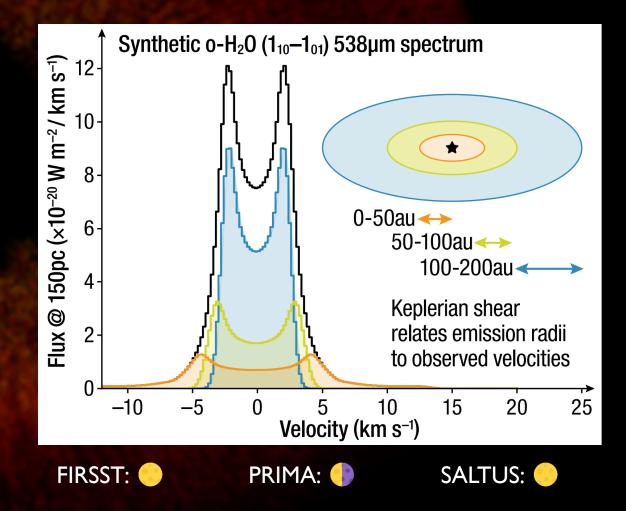
Water lines in prestellar cores and disks

- Complex emission and absorption features demand R=10⁶ to accurately measure line fluxes
- Can study kinematic behavior like infall to connect to later stages of planet formation
- Keplerian line profiles of disk water lines can inform us about the radial distribution of water
- Measurements of HDO/H_2O and $O/P H_2O$ ratios of prestellar cores and disks will clarify inheritance



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Debris disk science

 Yet undetected amorphous and crystalline ice features at 40-55 μm

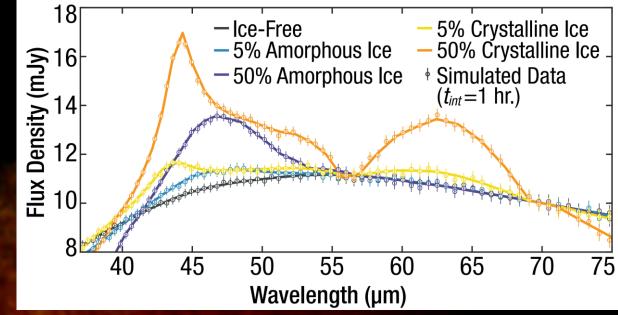
FIRSST: 🔵

PRIMA: SALTUS: O

- C:O ratios in gaseous debris disks can constrain composition and connect to exoplanet atmospheres
 - Need R>10⁴ for [CII] at 158 μm for total flux in most disks
- FIRSST: 🤚 PRI

PRIMA: 🌗

SALTUS: 😑



Other Science Objectives and GO Ideas

• Comets!

- D/H ratios from water lines
- Disk Evaporation Mechanisms
 [OI] and [NII]
- Revisit the Fab Four Debris Disks
 - Debris Disk Halos?
- Search for Exo-Zodis around nearby stars

- Extreme Debris Disk Variability
- Protostar Variability

• PRIMA GO Handbook

- Star-forming region mapping
- Size distribution of KBOs
- Mineralogy in PPDs
- Ammonia chemistry in PPDs

Discussion: What is missing on this list?

Thanks for listening!